Mem. S.A.It. Vol. 84, 38

Memorie della

Constraints on mass loss of globular clusters in dwarf galaxies

S. S. Larsen[1], J. Strader[2], and J. P. Brodie[3]

[1] Department of Astrophysics / IMAPP, Radboud University Nijmegen, P.O. Box 9010, 6500 GL Nijmegen, The Netherlands, e-mail: s.larsen@astro.ru.nl
[2] Department of Physics and Astronomy, Michigan State University, East Lansing, Michigan 48824, USA
[3] UCO/Lick Observatory, University of California, Santa Cruz, CA 95064, USA

Abstract. The Fornax dwarf spheroidal galaxy is well known for its very high globular cluster specific frequency, $S_N \approx 26$. Furthermore, while the field star metallicity distribution peaks at [Fe/H] ≈ -1, four of the five GCs have [Fe/H] < -2. Only about 5% of the field stars have such low metallicities. Hence, a very large fraction of about 1/5-1/4 of the most metal-poor stars belong to the four most metal-poor GCs. This implies that these clusters could, at most, have been a factor of $\sim 4-5$ more massive initially. A second, even more extreme case may be the IKN dwarf galaxy where $S_N \approx 124$. Although metallicities are not accurately known, the GCs account for about 13% of the *total* V-band luminosity of IKN.

Key words. Galaxies: individual: Fornax dSph – Galaxies: star clusters: general – Galaxies: stellar content

1. Introduction

One of the major puzzles related to the phenomenon of multiple stellar populations in globular clusters (GCs) is the "mass budget problem". Half or more of the present-day stellar mass in GCs typically resides in the 2nd generation, whose anomalous chemical abundances suggest that it formed out of ejecta produced by the first generation of stars. In most scenarios, this requires that the first generation was initially much more populous than it is today, by perhaps an order of magnitude or more (e.g. Gratton, Carretta & Bragaglia 2012). A test of whether such scenarios are viable would therefore be to determine how many stars in a galaxy can be traced back to GCs. In the Milky Way, this may be quite difficult in practice (but perhaps worth contemplating e.g. via chemical tagging; Bland-Hawthorn et al. 2010). However, dwarf galaxies may hold promising potential for carrying out this test. Although their GC populations are relatively modest in absolute terms, the few GCs that are present in some dwarfs can account for a substantial fraction of the total number of stars in the galaxy. Even without detailed constraints on the chemical composition of the stars within and out of the GCs, the ratio of GCs to field stars in a suitably restricted metallicity interval may thus be sufficiently high to provide useful constraints on how much mass the clusters could have lost.

Send offprint requests to: S. S. Larsen

The GC specific frequency, S_N[1], varies substantially among galaxies, with spiral galaxies typically having $S_N \approx 1$ and elliptical galaxies having $S_N \approx 3$–5. Even richer GC systems are found around cD galaxies, which can have S_N up to 10–15. The highest GC specific frequencies are found in dwarf galaxies, of which the Fornax dSph is a well-known example. For $M_V = -13.2$ (Mateo 1998), its five GCs (Hodge 1961) correspond to $S_N = 26$. Fornax does not appear to be an outlier or even particularly extreme in this respect; other examples of dwarf galaxies with very high specific frequencies are listed in Peng et al. (2008) and Georgiev et al. (2010).

These global numbers hide the fact that individual stellar sub-populations within galaxies can have much higher GC specific frequencies. In the Milky Way, about 2/3 of the 150 ancient GCs currently known (Harris 1996) are associated kinematically, chemically and spatially with the halo, although this component only accounts for about 1% of the stellar mass in our Galaxy. This "second specific frequency" problem is not restricted to our Galaxy; also in elliptical galaxies a disproportionately large fraction of the GCs appear to be associated with the metal-poor stars (Forte, Strom & Strom 1981; Harris & Harris 2002). However, in the Milky Way halo, the GCs still only *currently* account for ∼ 2% of the stellar mass (Kruijssen & Portegies Zwart 2008). This fraction may, however, have been very much higher in the past.

2. Abundances of stars and GCs in Fornax

The Fornax dSph exhibits the second specific frequency problem. The field star metallicity distribution has a broad peak around [Fe/H] ≈ −1 (Battaglia et al. 2006; Kirby et al. 2011) with only a small fraction of the stars having [Fe/H] < −2. In contrast, four of the five GCs have low metallicities around [Fe/H] ≈ −2. The specific frequency of the GCs, relative to field stars of the same metallicity, must therefore be very high indeed.

Table 1. Iron abundances and [Ca/Fe] ratios for the five GCs in the Fornax dSph from high-dispersion spectroscopy.

	[Fe/H]	[Ca/Fe]	Source
Fornax 1	−2.5 ± 0.1	+0.15 ± 0.04	L06
Fornax 2	−2.1 ± 0.1	+0.20 ± 0.03	L06
Fornax 3	−2.3 ± 0.1	+0.25 ± 0.08	L12
Fornax 4	−1.4 ± 0.1	+0.13 ± 0.07	L12
Fornax 5	−2.1 ± 0.1	+0.27 ± 0.09	L12

Accurate metallicity determinations are now available both for field stars and all the GCs in the Fornax dSph, allowing a more detailed analysis. Battaglia et al. (2006) published measurements of the Ca II IR triplet for 562 field red giants, covering the full radial range (∼ 1°). A comparison with [Fe/H] abundances derived from high-dispersion spectroscopy for a subset of the stars showed that the Ca II triplet measurements are reliable proxies for the Fe abundances (within 0.1–0.2 dex) for [Fe/H] > −2.5 (Battaglia et al. 2007). Measurements of [Fe/H] abundances by direct spectral fitting (Kirby et al. 2011) yield a very similar metallicity distribution to that derived by Battaglia et al. (2006).

Abundance measurements from high-dispersion spectroscopy of individual stars are available for three of the GCs (Fornax 1, 2 and 3) (Letarte et al. 2006, L06). We have recently carried out a detailed abundance analysis based on high-dispersion spectroscopy of *integrated* light for Fornax 3, 4 and 5 (Larsen, Brodie & Strader 2012, L12). For Fornax 3 our [Fe/H] abundance agreed with that measured by Letarte et al. within 0.1 dex. For Fornax 5 we found a low metallicity ([Fe/H]=−2.1) while Fornax 4 was confirmed to be more metal-rich than the other four clusters ([Fe/H]=−1.4 ± 0.1). The high-dispersion spectroscopic Fe and Ca abundances are summarized in Table 1. All five GCs have somewhat super-solar [Ca/Fe] abundance ratios, in agreement with measurements of

[1] $S_N = N_{GC} \times 10^{0.4(M_V-15)}$ where N_{GC} is the number of GCs in a galaxy and M_V the absolute V magnitude of the galaxy (Harris & van den Bergh 1981)

field star abundances for Fornax and other dSphs (Tolstoy, Hill & Tosi 2009).

3. Stars in the field vs. GCs

We corrected the Battaglia et al. metallicity distribution for spatial coverage and also took into account that the relative number of RGB stars above the spectroscopic magnitude limit depends on age and metallicity. Details are in Larsen, Strader & Brodie (2012). Figure 1 shows the global metallicity distribution of the field stars, corrected for these selection effects. Also shown is the metallicity distribution of the GCs, where each GC has been counted as $6 \times N_{\rm stars} \times 10^{-0.4(M_{V,\rm GC}+13.2)}$. Here, $M_V = -13.2$ is the M_V magnitude of the Fornax dSph, $M_{V,\rm GC}$ are the M_V magnitudes of the individual GCs, and $N_{\rm stars} = 562$ is the number of stars in the Battaglia et al. sample. Thus, the scale of the GCs is exaggerated by about a factor of 6.

It is immediately obvious from Fig. 1 that the metallicity distribution of the GCs differs enormously from that of the field stars.

As a first estimate of the fraction of metal-poor stars associated with the GCs, we simply scale the total luminosity of the Fornax dSph by the fraction of stars with [Fe/H] < -2. With the corrections mentioned above, this fraction is 5%, so that the integrated absolute magnitude of these stars is $M_V = -10.0$. The integrated magnitude of the four metal-poor GCs is $M_{V,\rm GCs} = -8.9$ (Webbink 1985). In other words, *the four metal-poor GCs account for more than 1/4 of the luminosity of all stars with metallicities* [Fe/H] < -2. A correction for the age- and metallicity dependent mass-to-light ratio of the field stars tends to further increase this fraction by 10%–20% (Larsen, Strader & Brodie 2012). This comparison is largely independent of assumptions about the IMF.

We can also estimate the fraction of stars by *mass* that reside in the metal-poor GCs, although this involves more assumptions. Recent estimates of the stellar mass of the Fornax dSph range from $4.3 \times 10^7 M_\odot$ to $6.1 \times 10^7 M_\odot$ (de Boer et al. 2012; Coleman & de Jong 2008). The mean virial M/L_V ratio of the metal-poor GCs is estimated to be $M/L_V \approx 3.5$, corresponding to a total mass of $1.0 \times 10^6 M_\odot$ for these GCs (Larsen, Brodie & Strader 2012). After correcting for age- and metallicity effects, and conservatively adopting the larger of the above estimates for the stellar mass of Fornax, we again find that a fraction of 1/5-1/4 of the metal-poor stars in Fornax are associated with the four metal-poor GCs. This comparison is, however, complicated by the relatively high observed M/L_V ratio of the GCs; a lower value of $M/L_V \approx 2$ is predicted for a standard Kroupa- or Chabrier-type IMF at these metallicities. If we were to adopt this lower M/L_V ratio for the GCs (albeit at odds with the dynamical measurements), the mass fraction of metal-poor stars in the GCs would decrease to 1/9-1/7.

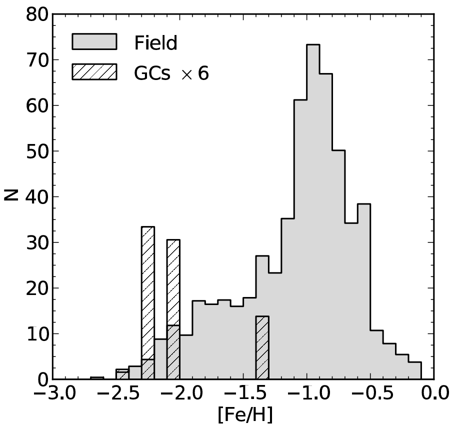

Fig. 1. Distribution of [Fe/H] abundances for field stars and GCs in the Fornax dSph.

4. Dwarf galaxies and multiple populations in GCs

Regardless of the uncertainties in the above analysis, it is clear that a very large fraction of the metal-poor stars in the Fornax dSph are *currently* – even after a Hubble time – members of GCs. The implication is that the Fornax GCs could not have been much more than a factor of 4–5 more massive initially – there simply aren't enough metal-poor stars present in the galaxy today to account for any more mass loss from the GCs. Furthermore, this must be considered a conservative upper limit as it assumes that no other stars with [Fe/H] < -2 formed in the field, or in other clusters that have since dissolved.

One potential caveat is that stars may have been lost from the Fornax dSph over its lifetime. However, Peñarrubia et al. (2009) studied the spatial profiles of several dSph galaxies in the Local Group, and concluded, based on N-body simulations of the signatures of tidal stripping that "The Fornax dwarf provides probably the clearest case of a dwarf whose stellar component has *not* been disturbed by Galactic tides" (their emphasis).

The issue would become moot if the Fornax GCs do not host the multiple populations that appear to be so ubiquitous in Galactic GCs. This, by itself, would of course be interesting – it would imply that GC formation in dwarf galaxies is quite a different process from that in larger galaxies (or, at least, the Milky Way). It is therefore worthwhile examining the constraints on the presence of multiple stellar populations in the Fornax GCs. The challenge here is that the clusters are too far away to easily observe large numbers of individual member stars spectroscopically. However, in their observations of 3 RGB stars in each of the clusters Fornax 1, 2 and 3, Letarte et al. (2006) found at least one star (in Fornax 3) which appears to have depleted Mg and O and enhanced Na, suggesting that the chemical abundance anomalies are present at least in this cluster. In our analysis of the integrated-light spectra, we found Mg to depleted relative to Ca and Ti in Fornax 3, 4 and 5, possibly a hint of the Mg-Al anticorrelation. While these data suggest that the Fornax GCs are not very different from their Milky Way counterparts, stronger constraints on the presence (or absence) of multiple stellar populations are clearly desirable.

Another question is whether other cases like Fornax can be found. One such case may be the IKN dwarf spheroidal in the Ursa Major group. This galaxy is even fainter than Fornax ($M_V = -11.5$; Georgiev et al. 2010) but also hosts 5 GCs. This makes the GC specific frequency of IKN even higher than that of Fornax, at a staggering $S_N = 124$. Photometry of RGB stars in IKN reveals a broad metallicity distribution reminiscent of that in Fornax (Lianou, Grebel & Koch 2010). The metallicities of the GCs are not accurately known, but their blue integrated colours suggest relatively low metallicities (Georgiev et al. 2010). However, without even trying to match the metallicities of GCs and field stars, the GCs account for about 13% of the *total* integrated V-band luminosity of the IKN galaxy.

Acknowledgements. SSL is grateful to the Leids Kerkhoven-Bosscha Fonds for supporting his attendance of this conference by means of a travel grant.

References

Battaglia, G. et al. 2006, A&A, 459, 423
Battaglia, G. et al. 2007, MNRAS, 383, 183
Bland-Hawthorn, J. et al. 2010, ApJ, 721, 582
Coleman, M. G., & de Jong, J. T. A. 2008, ApJ, 685, 933
de Boer, T. J. L. et al. 2012, preprint `arXiv:1206.6968`
Forte, J. C., Strom, S. E., & Strom, K. M. 1981, ApJ, 245, L9
Georgiev, I., et al. 2010, MNRAS, 406, 1967
Gratton, R. G., Carretta, E., & Bragaglia, A. 2012, A&AR, 20, 1
Harris, W. E. 1996, AJ 112, 1487
Harris, W., & Harris, G. 2002, AJ, 123, 3108
Harris, W. E., & van den Bergh, S. 1981, AJ, 86, 1627
Hodge, P. 1961, AJ, 66, 83
Kirby, E. et al. 2011, ApJ, 727, 78
Kruijssen, J. M. D., & Portegies Zwart, S. F. 2008, ApJ, 698, L158
Larsen, S. S., Brodie, J. P., & Strader, S. 2012, A&A, 546, A53
Larsen, S. S., Strader, J., & Brodie, J. P. 2012, A&A, 544, L14
Letarte, B. et al. 2006, A&A 453, 547
Lianou, S., Grebel, E. K., & Koch, A., 2010, A&A, 521, A43
Mateo, M. 1998, ARA&A, 36, 435
Peñarrubia, J., et al. 2009, ApJ, 698, 222
Peng, E. W., et al. 2008, ApJ, 681, 197
Tolstoy, E., Hill, V., & Tosi, M. 2009, ARA&A, 47, 371
Webbink, R. F. 1985, in: IAU Symp. 113, 541